

Comparison of steady-state flow calculations with M1 and FLD

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What is PNS and PNSC simulation

Proto-Neutron Star (**PNS**) is an object produced by a **core collapse supernova** and it is **cooled** by the **emission of neutrinos** (PNS cooling, **PNSC**), which evolves into the Neutron Star (NS).

This PNSC simulation requires solving the **Einstein equation** (spacetime), the **hydrodynamic equation** (other than neutrinos) and the **Boltzmann equation** (neutrinos).

Due to the **multidimensionality** of the Boltzmann equation, this equation cannot be computed for long periods of time without **the use of approximations**.

Numerical cost is too high!

1. Objective and Background

- ✓ A code describing the **cooling process of proto-neutron stars** (PNSC) was created by H.Suzuki (1993) using multi-energy flux limited diffusion (FLD) scheme to solve Boltzmann equation approximately.
- ✓ To prepare for **future neutrino observations**, the current code needs to be improved and updated to make it **more accurate** and capable of **long-time** calculations.
- ✓ Therefore, taking these points into account, I created a code to solve the Boltzmann equation by **M1** and compared it to FLD.

2. Formalism

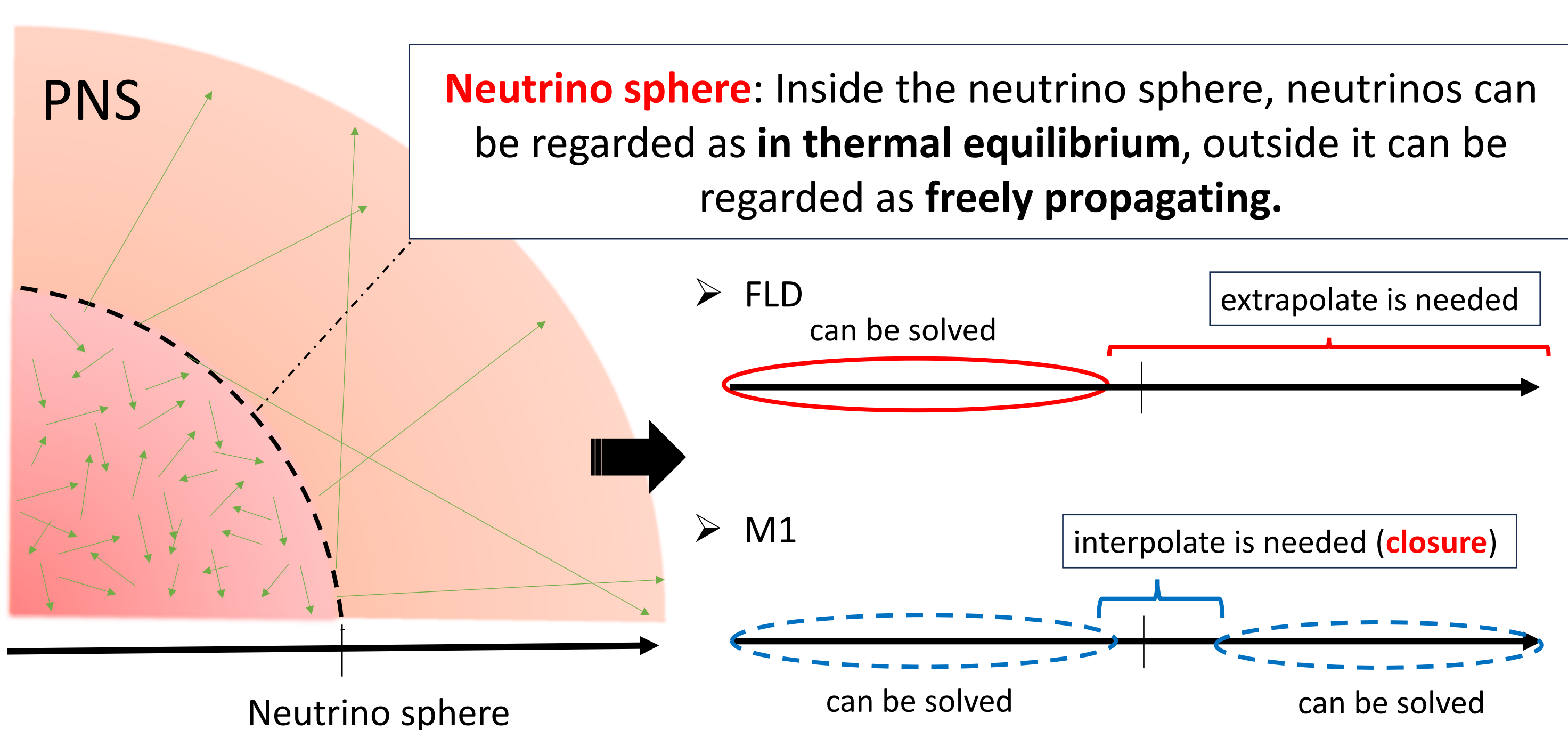
Difference between FLD and M1

✓ Flux limited diffusion (original code)

Assume diffusion limit of neutrinos (good assumption inside neutrino sphere) and set a limit of the flux in transparent region.

✓ M1-closure (Created Code)

The two states of the neutrino, **thermal equilibrium** and **free propagation**, can be solved exactly. The states in between are complementary (closure).



Equation for moments (Steady-flow for simplicity)

$$\partial_t \vec{u} + \partial_m (4\pi r^2 \rho \vec{F}) = \vec{s}$$

f_ν : distribution function of neutrino

Moment

Variables

$$n_\nu = \frac{\omega^2}{(hc)^3} \int f_\nu d\Omega$$

$$\vec{u} = \frac{1}{\rho} (n_\nu, F_\nu)$$

$$\vec{F} = \frac{1}{\rho} (F_\nu, P_\nu)$$

$$F_\nu = \frac{\omega^2}{(hc)^3} \int f_\nu \mu d\Omega$$

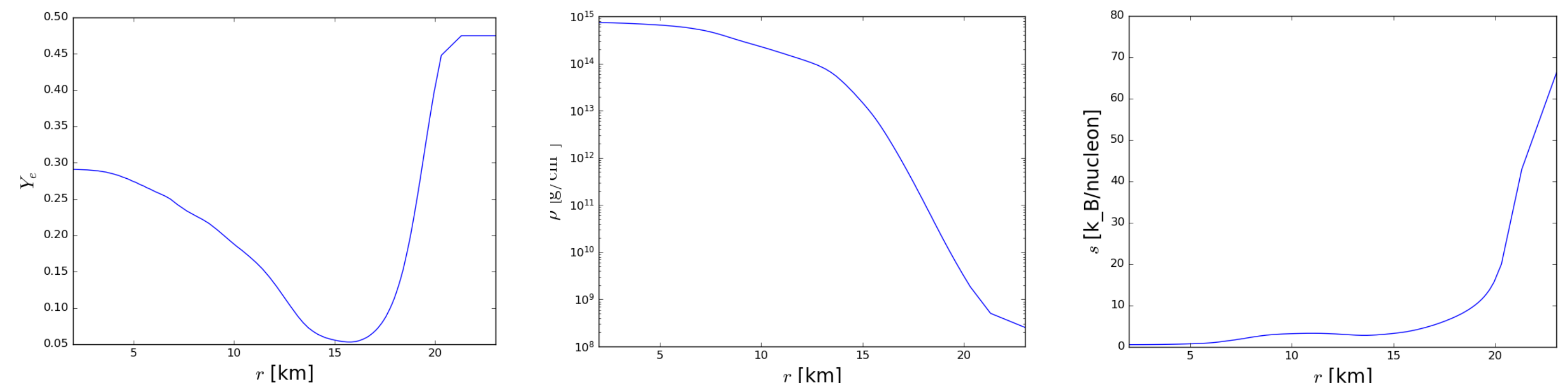
$$\vec{s} = \frac{1}{\rho} (S^0, S^1 - 4\pi r c \rho e^\phi (n_\nu - P_\nu) \partial_m (r e^{-\phi}))$$

$$P_\nu = \frac{\omega^2}{(hc)^3} \int f_\nu \mu^2 d\Omega$$

Closure : Levermore, ME, etc.

3. Numerical situation

- ✓ To do the numerical simulations of steady flow, I imposed the following situation.
- ✓ These profile obtained from the result of the PNSC calculation (original code).
- after 600 msec of shock revival



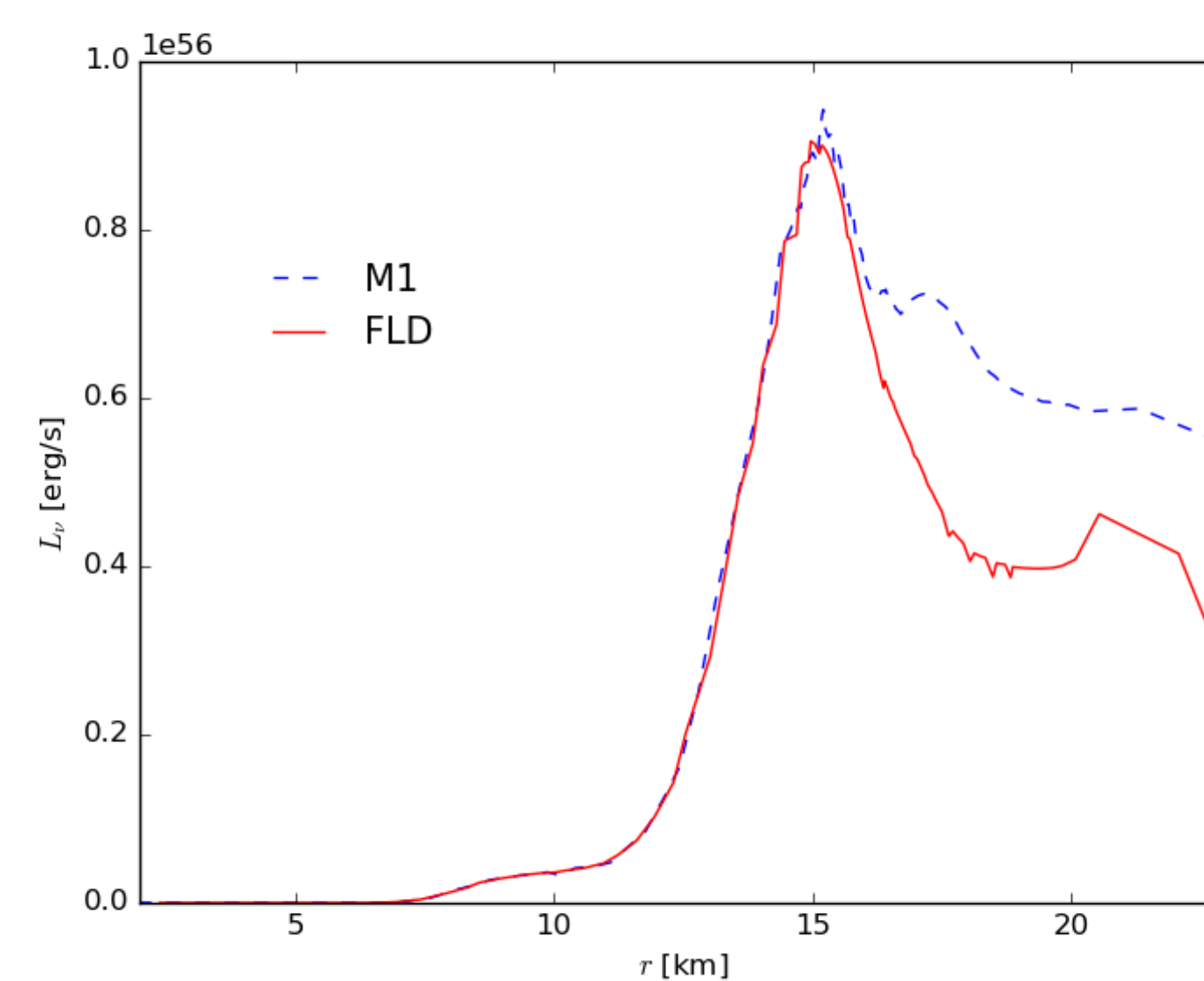
4. Calculation results

- ✓ Compare M1 and FLD in steady flow with a fixed distribution of fluid after 600 msec of proto-neutron star cooling.
- ✓ Compare the results of different closures.

Details

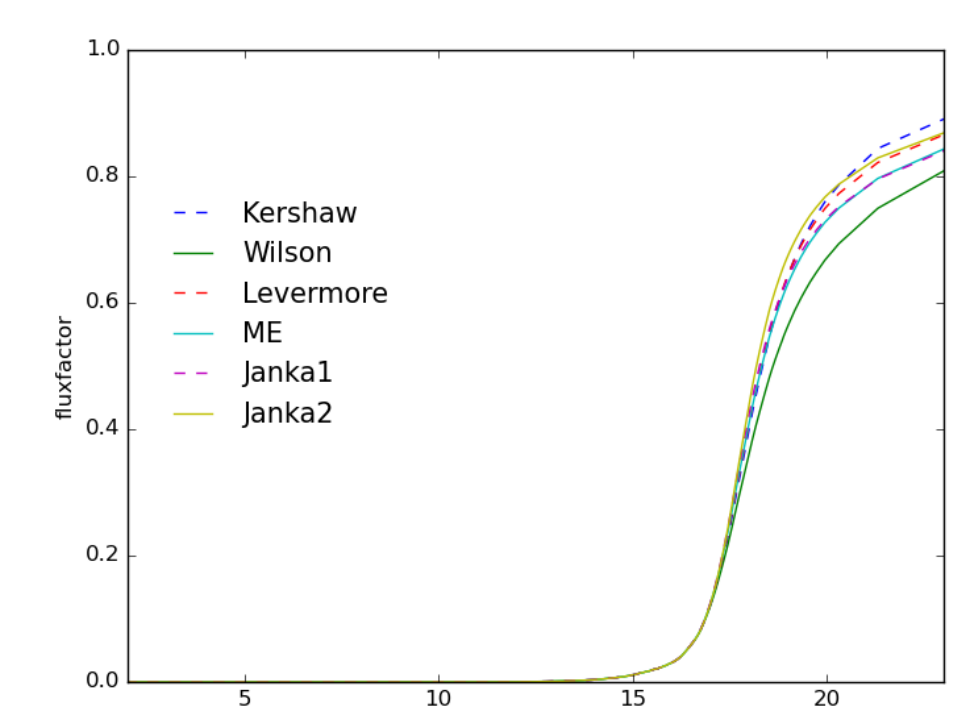
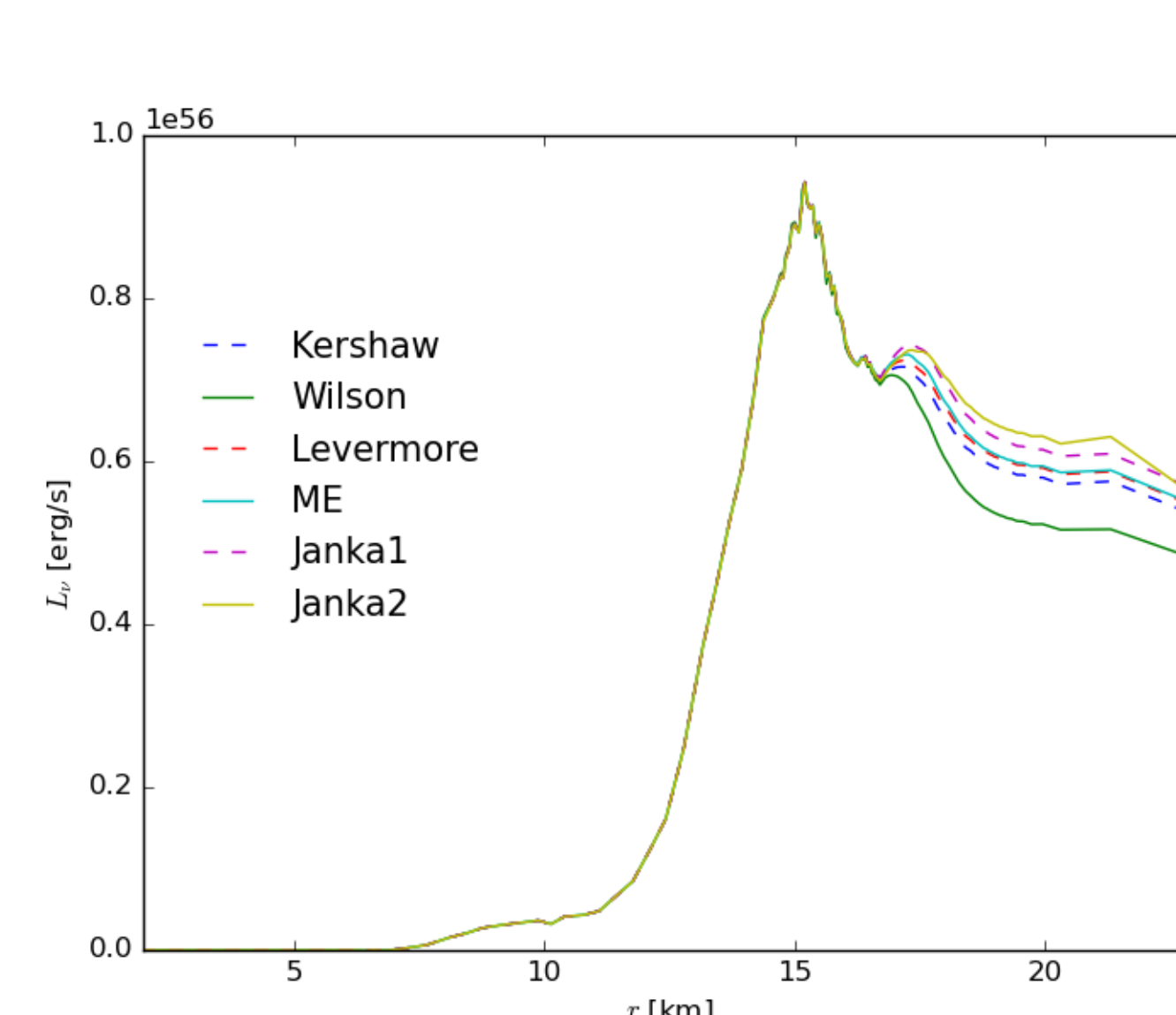
I. Compare M1 and FLD

Luminosity ($\frac{dL_\nu}{dE_\nu} \Delta E_\nu$, observed as a neutrino of 20 MeV at infinity)



- Near the center, the neutrino is in **thermal equilibrium**, so the results for FLD and M1 are the same.
- The difference between FLD and M1 appears **near the neutrino sphere** (~17 km).
- The higher the energy, the larger the neutrino sphere and the smaller the difference between FLD and M1.

II. Compare with different closures



- Each closure behaves differently **near the surface**.
- It is **impossible** to know which closure is better without comparing it to the more accurate solution.

5. Conclusion/Future Plans

- Neutrino luminosity were found to be larger for M1 than for FLD.
- However, the result for M1 is also an **approximation**, and the actual value may be somewhere between M1 and FLD.
- Therefore, we would like to discuss the behavior of M1 by **comparing** the more accurate or direct integration solution of the Boltzmann equation (of course, due to numerical cost, long time calculation is impossible) with M1 in the future.